

# Diagnostic of the temperature and differential emission measure (DEM) based on *Hinode*/XRT data

M. Siarkowski<sup>1</sup>, R. Falewicz<sup>2</sup>, A. Kępa<sup>1</sup>, and P. Rudawy<sup>2</sup>

<sup>1</sup>Space Research Centre, Polish Academy of Sciences, 51-622 Wrocław, ul. Kopernika 11, Poland

<sup>2</sup>Astronomical Institute, University of Wrocław, 51-622 Wrocław, ul. Kopernika 11, Poland

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**Abstract.** We discuss here various methodologies and an optimal strategy of the temperature and emission measure diagnostics based on *Hinode* X-Ray Telescope data. As an example of our results we present the determination of the temperature distribution of the X-rays emitting plasma using a filters ratio method and three various methods of the calculation of the differential emission measure (DEM). We have found that all these methods give results similar to the two filters ratio method. Additionally, all methods of the DEM calculation gave similar solutions. We can state that the majority of the pairs of the *Hinode* filters allows one to derive the temperature and emission measure in the isothermal plasma approximation using standard diagnostics based on the two filters ratio method. In cases of strong flares one can also expect good conformity of the results obtained using a Withbroe – Sylwester, genetic algorithm and least-squares methods of the DEM evaluation.

**Keywords.** Solar physics, astrophysics, and astronomy (Corona and transition region; Ultraviolet emissions; X rays and gamma rays; Instruments and techniques)

## 1 Introduction

The solar corona with temperatures greater than one million degrees emits mainly in the X-ray domain of the spectrum. Its emission contains both the continuum from thermal plasma and the line emission of highly ionized elements. The earliest phases of solar flares are dominated by non-thermal emissions (hard X-rays), the remainder of the events manifests itself primarily by thermal emissions at wavelengths ranging from X-rays through visual. The thermal radiation comes from plasma having a wide range of temperatures.

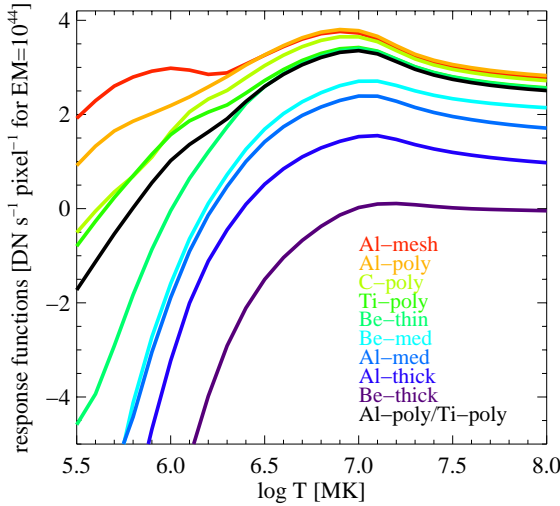
Temperature and emission measure diagnostics of the active region and flaring loops are very important to model the

loop structures, and analyze the heating sources and their (heating) distribution along the loop.

For the optically thin emission, assuming isothermal plasma temperature, the well known method of filter ratio, based on the ratio of the observed fluxes in two selected energy (or wavelengths) bands (Vaiana et al., 1973; Thomas et al., 1985) can be used. In the case of multi-temperature plasma the filter ratio method derives the line-of-sight average isothermal temperature and emission measure. Therefore, in this approach it is more adequate to use many filters or lines in terms of differential emission measure distributions  $DEM(T)/dT$  (Pottash et al., 1964; Sylwester et al., 1980). Extended discussion of the DEM versus isothermal approximation can be found, for example, in Schmelz et al. (2001), Cirtain et al. (2007) and references therein.

*Hinode* is a new advanced optical, UV and X-ray solar satellite launched on 22 September 2006. It is equipped with three instruments: Solar Optical Telescope (SOT), EUV Imaging Spectrometer (EIS) and the X-Ray Telescope (XRT). XRT is a high resolution Wolter-I type grazing incidence telescope, which is sensible to the emission of plasma in the temperature range  $5.5 < \log T < 8$ . Such wide temperature coverage is realized by nine X-ray filters, each with its own passband and therefore a different response to plasma temperature. These filters are mounted on two filter wheels which allow many combinations of filters and thus can be used for a temperature diagnostic. Figure 1 presents the theoretical temperature response functions of the ten XRT filters combination for a unit volume emission measure of  $10^{44} \text{ cm}^{-3}$ , for each filter used in this paper. XRT has  $2048 \times 2048$  back-illuminated CCD array with a  $13.5 \mu\text{m}$  pixel size. It allows the Sun soft X-ray images to make a  $34 \times 34$  arcmin FOV with a spatial resolution of  $\sim 1$  arcsec. Exposure times are in the range from 1 ms to 16 s. A detailed description of the instrument can be found in the paper by Golub et al. (2007).

Correspondence to: M. Siarkowski  
(ms@cbk.pan.wroc.pl)



**Fig. 1.** Theoretical temperature response functions of the selected XRT filters.

In this paper we will check if the XRT response characteristics allow diagnostics of temperature ( $T$ ) and emission measure ( $EM$ ) of the coronal plasma. We also describe the observations selected for the analysis. Additionally, we present and discuss the filters ratio method in isothermal approximation and three different methods for differential emission measure (DEM) analysis in multi-temperature approximation, respectively.

In all calculations we used the temperature response function as given by SolarSoft and presented in Fig. 1.

## 2 Data analysis

For our work we selected three typical X-ray emitting active regions observed by *Hinode*/XRT with 6 or more different filters. To prepare the selected data we used standard SolarSoft routine *xrt\_prep.pro*, which is similar in nature to *sxt\_prep.pro* (Yohkoh software). This routine converts raw data into index and structure arrays. Additionally, this process removes several instrumental effects. We could mention some of the most important, for example: replacement of the near-saturated pixels for a value greater than 2500 DN with 2500 DN, remove radiation-belt/cosmic-ray and streaks, calibration for read-out signals, removing of the CCD bias, calibration of the dark current, normalization of each image for exposure time. Using processed data we calculated their differential emission measure distributions and using a two filters ratio method, temperature and emission measure maps in isothermal approximation.

### 2.1 Two filters ratio diagnostic

The filter-ratio method is a technique commonly used for plasma diagnostics based on data from broad-band X-ray imaging systems. This method requires analysis of two or more intensity maps (images) of the investigated structures obtained with various broad-band filters. It delivers maps of temperatures and line-of-sight emission integrals or emission measure distributions as a function of temperature.

For a multi-band instruments like XRT, under the assumption that the emitting plasma is optically thin, the observed flux  $F_k$  in filter  $k$  in a single pixel is:

$$F_k = \int_T \int_\lambda f(T, \lambda) \eta_k(\lambda) EM(T) d\lambda dT, \quad (1)$$

where  $f(T, \lambda)$  is coronal plasma emissivity as a function of the plasma temperature  $T$  and the wavelength  $\lambda$ , calculated at a distance of 1 AU using a CHIANTI spectral code (including both emission lines and continuum).  $\eta_k(\lambda)$  is an effective area of the instrument which includes the geometric area and the reflectance of the telescope mirror, the entrance and the focal plane filter transmissions and the CCD quantum efficiency.  $EM(T)$  is a collection of emission measures over a range of temperatures located along the line of sight. Defining the temperature response function  $P_k(T)$  of filter  $k$  as a convolution of the instrument effective area  $\eta_k$  and emissivity of solar plasma  $f$ :

$$P_k = \int_\lambda f(T, \lambda) \eta_k(\lambda) d\lambda, \quad (2)$$

we can rewrite Eq. (1) as:

$$F_k = \int_T P_k(T) EM(T) dT \approx \sum_i P_k(T_i) EM(T_i). \quad (3)$$

For isothermal approximation this equation can be written as:

$$F_k = P_k(T) EM(T). \quad (4)$$

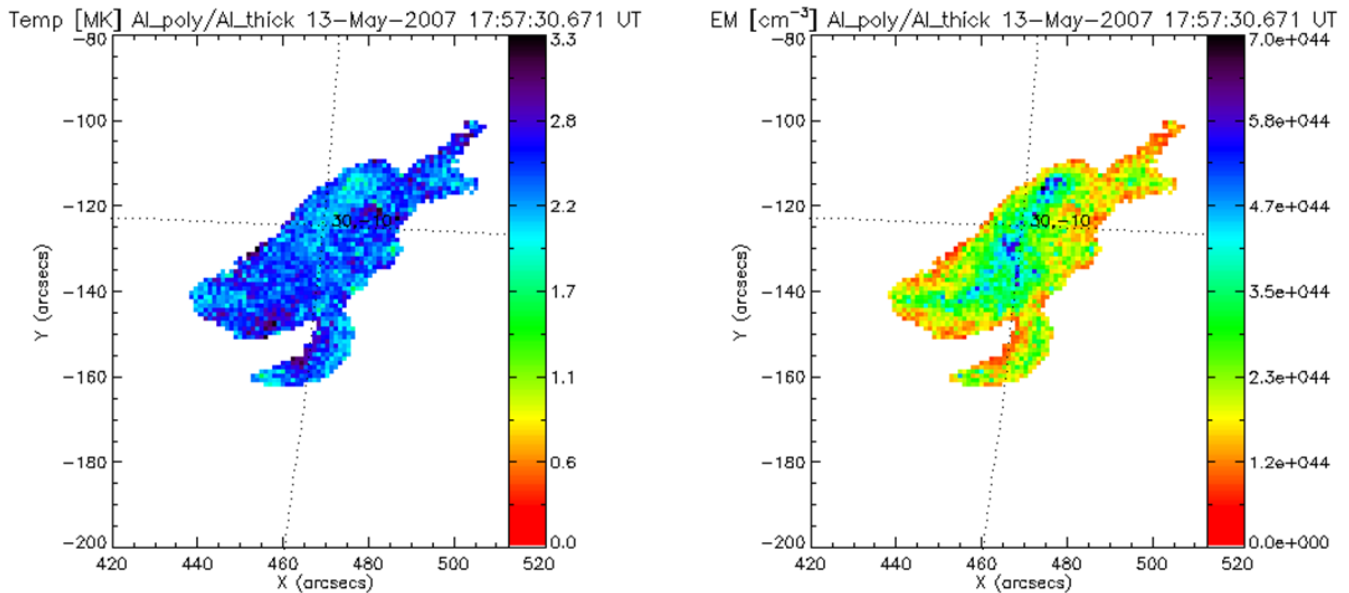
Similar to Thomas et al. (1985) using Eq. (4) one can calculate ratio  $R$  of the observed fluxes ( $F$ ) in two filters labeled “1” and “2” for a co-aligned pixel area. This can be written as:

$$R(T) = \frac{F_1}{F_2} = \frac{P_1(T)}{P_2(T)}. \quad (5)$$

According to Eq. (5), the ratio  $R$  of observed fluxes depends on plasma temperature  $T$  only. In the next step, one can calculate the emission measure  $EM(T) = F_k / P_k(T)$ .

This procedure can be used on all combinations of the XRT filters, but it has a limited usability. In particular, ratio  $R$  can be sometimes an ambiguous function in the whole range of temperatures.

Using the procedure described above we made temperature and emission measure maps for an area limited by an isoline 0.3 of the brightest pixel on the image.



**Fig. 2.** Temperature and emission measure maps calculated from two filters ratio, AL<sub>poly</sub>/AL<sub>thick</sub> for the active region NOAA 0955 observed on 13 May 2007 at 17:57:30 UT.

## 2.2 Analysed events

The first analysed region was a relatively faint active region NOAA 0955. Observations were made on 13 May 2007 at around 18:00 UT. Images taken in a ten filters combination are available. Temperature and emission measure maps were calculated using two pairs of filters, AL<sub>poly</sub>/AL<sub>thick</sub> and Ti<sub>poly</sub>/AL<sub>thick</sub>. As an example the resulted maps for the AL<sub>poly</sub>/AL<sub>thick</sub> pair are presented in Fig. 2. For both pairs of filters the filter ratios are sensitive for temperatures from  $\log(T)=6.1$  up to  $\log(T)=7.7$ . However, for all structures of this quiet active region we obtained temperatures in the ranges 2.0–3.3 MK and 1.8–2.4 MK for the AL<sub>poly</sub>/AL<sub>thick</sub> and Ti<sub>poly</sub>/AL<sub>thick</sub> pairs, respectively. The range of the observed temperatures is typical for solar quiet active regions (see, e.g. Golub and Pasachoff, 1997).

An active region NOAA 0923 was seen on the edge of the West solar limb between 19 and 23 November 2006. Observations analyzed here were made on 20 November 2006 at around 21:16 UT. Images taken in six filters are available. Temperature and emission measure maps were calculated for two pairs of filters: AL<sub>poly</sub>/AL<sub>mesh</sub> and AL<sub>poly</sub>/Be<sub>thin</sub>. The resulting maps for the AL<sub>poly</sub>/Be<sub>thin</sub> filters pair are presented in Fig. 3. The highest temperatures in the range of 3.5 MK are observed in the brightest loop, seen just on the solar limb.

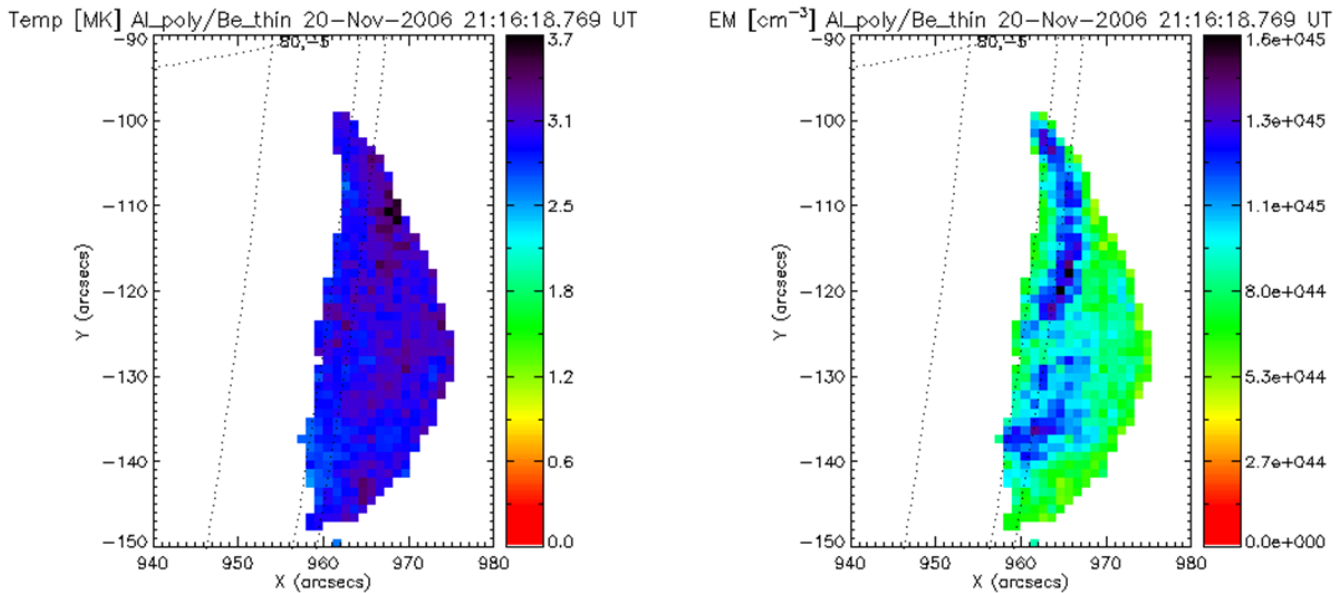
The same active region NOAA 0923 was observed one hour later with seven filters. Temperature and emission measure maps were calculated for two pairs of filters:

AL<sub>poly</sub>/AL<sub>mesh</sub> and AL<sub>poly</sub>/Be<sub>thick</sub>. The resulted maps for the AL<sub>poly</sub>/Be<sub>thick</sub> pair are shown in Fig. 4.

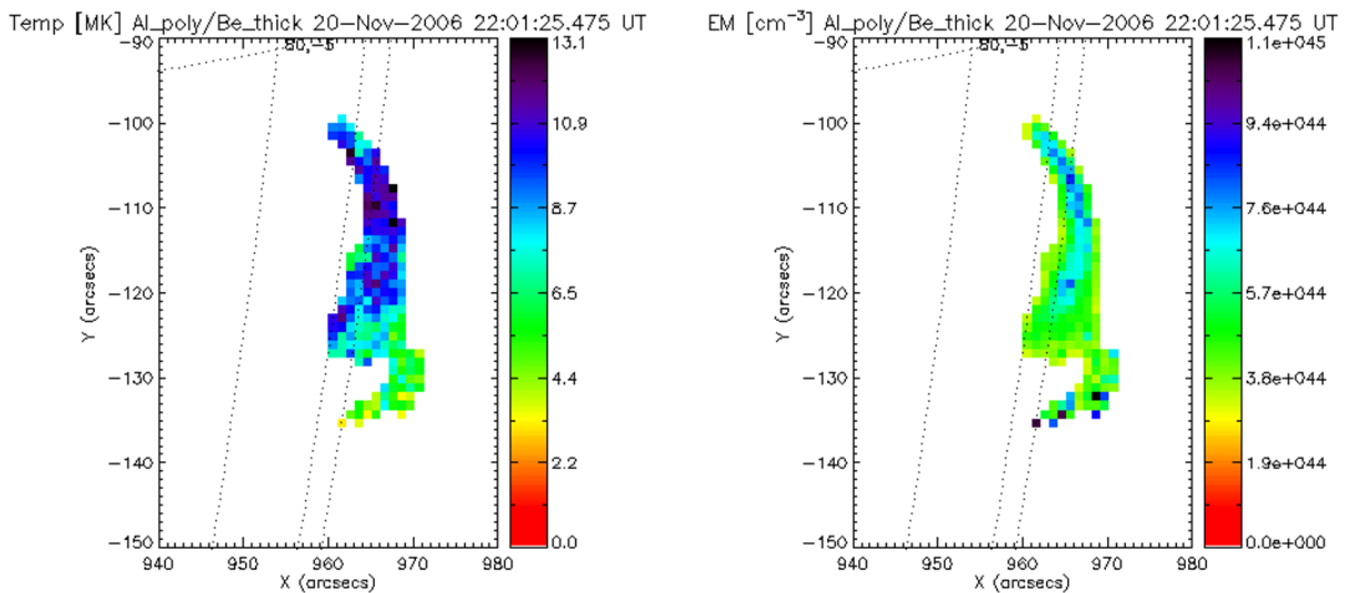
Due to the construction of the filter wheels, XRT images in various filters are only obtained consecutively. Thus an important assumption underlying the filters ratio diagnostic (and also DEM analysis described below) is that there was no substantial change in the morphology and energetics of the event between two images using various filters. This assumption is usually fulfilled in the case of quiet Sun and active region observations. Checking the light-curves for different filters we proved that this assumption is also true in the observations above described. Moreover, in the case of the small B1.5 flare images with different filters were obtained first around the maximum and the second one in decay phase of the event. We use this pair because the second filter had only one “thick” filter available. The exact co-alignment of the images is likewise important, especially for the DEM calculations in the individual pixels.

Using a “harder” AL<sub>poly</sub>/Be<sub>thick</sub> filters pair instead of a “softer” AL<sub>poly</sub>/Be<sub>thin</sub> pair results in obtaining much higher temperatures. Temperatures in the brightest loop reached average values up to  $\sim 13$  MK. Such a high temperature could be connected to the appearance of a small B1.5 class solar flare observed in this active region. The flare started at 22:00 UT, reached its maximum at 22:05 UT and ended at 22:09 UT. In fact, the temperature obtained from GOES data for this small flare reached the value around 10 MK.

It is possible with XRT to register images of reasonable S/N with cadences of 1 or 2 s. In the case of solar flares



**Fig. 3.** Temperature and emission measure maps calculated from two filters ratio, Al<sub>poly</sub>/Be<sub>thin</sub> for the active region NOAA 0923 observed on West solar limb on 20 November 2006 at 21:16:19 UT.

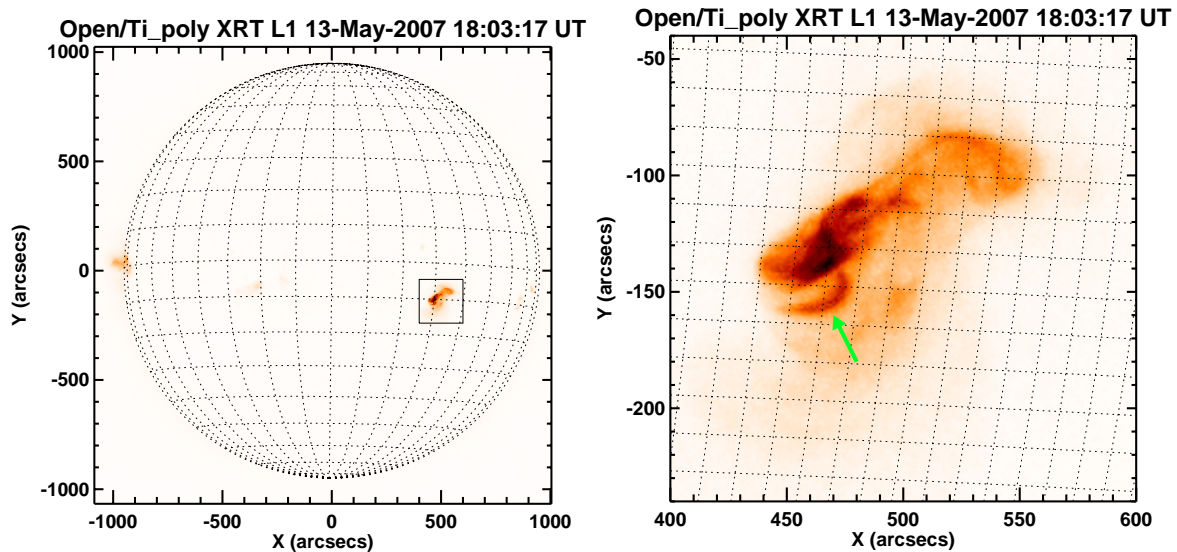


**Fig. 4.** Temperature and emission measure maps calculated from two filters ratio, Al<sub>poly</sub>/Be<sub>thick</sub> for the active region NOAA 0923 observed on West solar limb on 20 November 2006 at 22:01:25 UT.

where fluxes can vary rapidly it is possible, in principle, to interpolate the intensities of one image to match the time at which the image in the other filter was acquired. Examples of temperature and DEM analysis of solar flares are well known in the case of *Yohkoh* observations (McTiernan et al., 1999). Examples of DEM analysis with *Hinode*/XRT in the *GOES* C8.2 flare is presented by Reeves et al. (2007).

### 2.3 Distribution of the Differential Emission Measure

Differential emission measure, defined as  $DEM(T) = d(EM)/dT = n_e^2 dV/dT$ , is a useful diagnostic tool of non-isothermal plasma. It is usually calculated from a number of line fluxes formed at different temperature ranges. Generally the determination of the DEM distribution



**Fig. 5.** On the left: full Sun image obtained by XRT/*Hinode* with Ti\_poly filter. On the right: enlarged map of NOAA 955 active region observed by *Hinode*/XRT telescope with Ti\_poly filter. The DEM distribution, presented on Fig. 7, was calculated for the loop pointed by the green arrow.

is an ill-conditioned problem with no unique solution (Craig and Brown, 1976) and this problem needs some kind of regularization. Calculations of DEM from broad-band filters observations is also possible if a corresponding filters have different temperature response functions (see Fig. 1). In the case of the *Hinode* broad-band filter observation, such a method was worked out and tested by Weber et al. (2004).

In this paper we compare three different methods of the DEM reconstruction using *Hinode*/XRT observations with ten filters. The methods used are as follows:

**Genetic algorithm (GA)** – This algorithm is based on ideas of biological evolution and a natural selection mechanism. Starting from randomly chosen initial populations of different DEMs a new generation of DEMs is produced by crossover and mutations. The process of breeding (and multiplication) of the whole population is controlled by assumed fitness criterion, for example, DEMs with a higher fitness have a higher probability to participate in the process of multiplication using crossover. Examples of a DEM distribution calculated with the genetic algorithm can be found, for example, in papers by Kaastra et al. (1996), Guedel et al. (1997) and McIntosh et al. (2000).

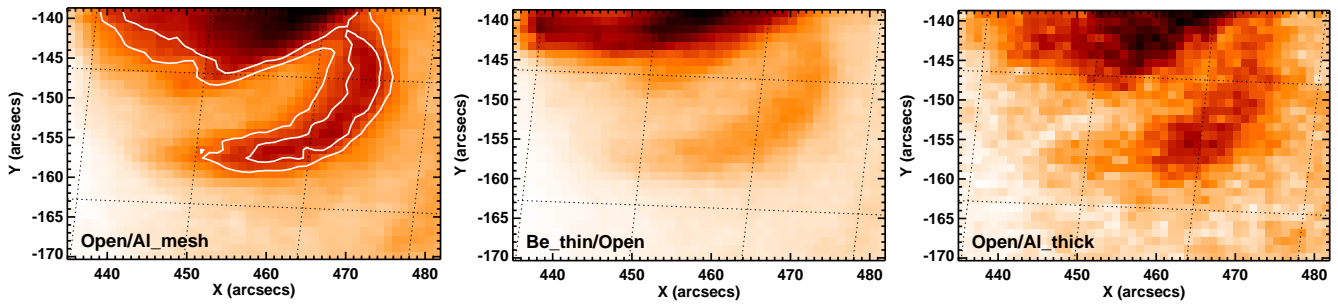
We used an IDL implementation of Charbonnaue’s “PIKAIA” Fortran code Charbonnaue (1995). Our population had 1000 individual DEM distributions or “genes”. After a process of multiplication a new population was created consisting of 90% new DEMs and 10% old DEMs with the highest fitness. We used  $\chi^2$  as the fitness criterion of the DEM. The process of evolution was stopped after 10 000–20 000 generations, when the convergence became very slow. The minimum values of  $\chi^2$  were around 4.3; typically 15

to 30 “best”, newly-generated DEMs had  $\chi^2$  values below 10. The evolution process was repeated 6 times, each time starting from a new random population. As a result we obtained 136 solutions with  $\chi^2 < 10$ . In Fig. 7 the green line presents the mean of these 136 solutions ( $\chi^2$  of the mean is 4.7). The errors were calculated as a standard deviation from the mean. The solution obtained with GA has  $\chi^2$  less than those obtained with WS and LS methods.

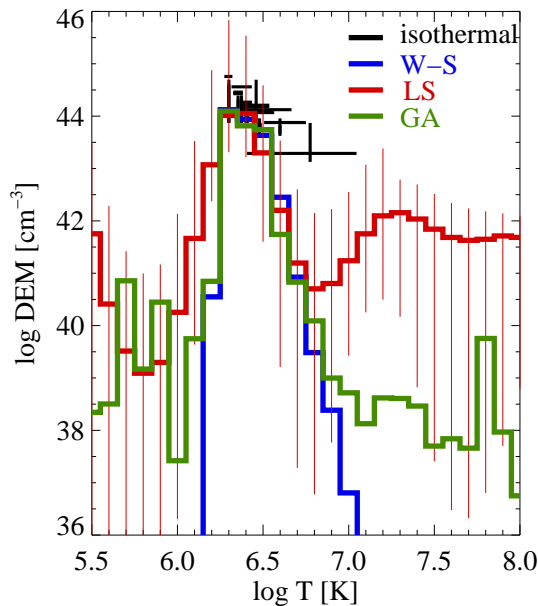
The Withbroe-Sylwester (WS) method is iterative, multiplicative algorithm relying on the maximum likelihood approach. In our calculations we started from a constant value DEM approximation. For detailed descriptions of this method, see Withbroe (1975) and Sylwester et al. (1980). Using the WS method we performed 200 simulations consisting of 3000 iterations each. The obtained  $\chi^2$  were in the range from 5 to 25. The mean solution has  $\chi^2 = 10$  (see Fig. 7).

**Least-squares (LS) method** – this procedure calculates DEM based on a least-squares fit of the calculated fluxes to the observed ones. The method is described in more detail by Weber et al. (2004) and implemented as a *xrt\_dem\_iterative* routine in SolarSoft. The DEM function is evaluated from some spline points, which are directly manipulated by the  $\chi^2$  fitting routine. Using the LS method we performed 200 LS-simulations. The obtained  $\chi^2$  were in the range from 20 to 30. The mean solution has  $\chi^2 = 25$ .

As an example we have calculated DEM for the loop observed on 13 May 2007 (see Figs. 5 and 6). The flux used in the DEM calculation was averaged over an area delimited by a 0.8 contour level on the Al\_mesh filter image (see Fig. 6, left panel).



**Fig. 6.** Images of the loop observed in active region NOAA 09555 around 18:00 UT, taken in three different *Hinode*/XRT filters. Contours overlotted on an Al<sub>mesh</sub> filter image corresponds to 0.6 and 0.8 of the maximum flux of the loop.



**Fig. 7.** The DEM distributions of the loop observed on 13 May 2007 obtained using the three analysed numerical methods and using isothermal approximation with different filter pairs.

We have tested all three methods using a synthetic isothermal model. The reconstructed width of the solutions allow us to estimate the temperature resolutions of the method. This resolution is low at the edges of the considered temperature range ( $5.5 < \log T < 8$ ) and increase in the temperatures where the response functions have their maxima (see Fig. 1), with best values of  $d\log T = 0.5\text{--}0.6$ ,  $0.2\text{--}0.3$  and  $0.1\text{--}0.2$  for the LS, WS and GA methods, respectively. DEM errors are similar for all three methods, and we show in Fig. 7 only error bars for LS, for readiness and transparency.

The uncertainties are in fact very large at lower temperatures (see Fig. 7), and for this reason all three methods give solutions lying within error limits of both other methods. Thus one cannot evaluate precisely *EM* in these temper-

atures using *Hinode*/XRT data. However, it is well known that DEM has a minimum around  $\log T \approx 5.5$  for quiet Sun and active regions, and for this reason, each obtained solution could be correct (Brosius et al., 1996; Schmelz et al., 2001).

There are also differences in DEM solutions obtained for high temperatures. The LS solution indicates the presence of a bigger amount of the plasma having the temperatures in the range ( $7.0 < \log T < 7.3$ ) than the WS and GA methods. The differences are larger than estimated errors. A possible explanation of these differences is that LS models depend on initial approximation of DEM (Weber et al., 2004). The LS method also has the lowest temperature resolution comparing to WS and GA methods. The difference between all three methods should be extensively tested using XRT observations.

Obtained  $\chi^2$  are rather high. Considering the whole loop instead of one or two pixels, the relative errors of flux decrease and thus  $\chi^2$  increase (being much above 1.0). On the other side, there are still large uncertainties in the fluxes measured with an Al<sub>thick</sub> filter which strongly influence DEM in high temperatures. A large  $\chi^2$  value can also be attributed to uncertainties in atomic data which contribute to the response functions and to a systematic error in background rejection.

### 3 Discussion and conclusions

XRT telescope can detect coronal plasma of various temperatures ranging from less than 1 MK up to more than 10 MK. By applying observations made with two or more filters one can analyze coronal temperature and emission measured along a line-of-sight.

The filter-ratio method has been used frequently in the past, but the high quality XRT data give temperature and emission measure maps with unprecedented accuracy and resolution. However, the filter-ratio method is not appropriate for multi-temperature structures, giving an average value of *T* and *EM* only. For such a case the DEM methods of

analysis of the distribution of emission measure should be used.

The DEM distributions obtained with the described methods are presented in Fig. 7. All three methods give the main peak of emission for  $\log T=6.3$ . Observed differences between results of the three methods lie more or less within the error limits (please note the logarithmic scale of DEM). This figure also presents a set of isothermal solutions obtained for different pairs of filters. Their values lie around the main peak of DEM obtained with numerical methods described before. Thus, we can conclude that all three methods give results similar to the two filters ratio method.

We can also conclude that described methods of the DEM calculation gave similar and comparable solutions. Additionally, based on the shapes of the emission functions (see Fig. 1) we can expect much better conformity of the results obtained with the WS, LS and GA methods in a case of strong flares. We can state that the majority of the pairs of the *Hinode* filters allows one to derive the temperature and emission measure in the isothermal plasma approximation using standard diagnostics based on the two filters ratio method. The temperatures evaluated with this method for two non-flaring active region plasmas were in the range from around 1.5 MK to about 5 MK.

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## References

- Brosius, J. W., Davila, J. M., Thomas, R. J., and Monsignori-Fossi, B. C.: Solar EUV spectroscopy with serts: measurements of active and quiet Sun properties, *ApJS*, 106, 143–164, 1996.
- Charbonneau, P.: Genetic Algorithms in Astronomy and Astrophysics, *ApJS*, 101, 309–334, 1995.
- Cirtain, J. W., Del Zanna, G., DeLuca, E. E., Mason, H. E., Martens, P. C. H., and Schmelz, J. T.: Active Region Loops: Temperature Measurements as a Function of Time from Joint TRACE and SOHO CDS Observations, *ApJ*, 655, 598–605, 2007.
- Craig, I. J. D. and Brown, J. C.: Fundamental limitations of X-ray spectra as diagnostics of plasma temperature structure, *A&A*, 49, 239–250, 1976.
- Golub, L., DeLuca, E., Austin, G., et al.: The X-Ray Telescope (XRT) for the *Hinode* Mission, *Solar Phys.*, 243, 63–86, 2007.
- Golub, L. and Pasachoff, J. M.: *The Solar Corona*, Cambridge: Cambridge Univ. Press, 101–104, 1997.
- Guedel, M., Guinan E. F., Mewe, R., Kaastra, J. S., and Skinner, S. L.: A Determination of the Coronal Emission Measure Distribution in the Young Solar Analog EK Draconis from ASCA/EUVE Spectra, *ApJ*, 479, 416–426, 1997.
- Kaastra, J. S., Mewe, R., Liedahl, D. A., Singh, K. P., White, N. E., and Drake, S. A.: Emission measure analysis methods: the corona of AR Lacertae revisited, *A&A*, 314, 547–557, 1996.
- McIntosh, S. W., Charbonneau, P., and Brown, J. C.: Preconditioning the Differential Emission Measure (Te) Inverse Problem, *ApJ*, 529, 1115–1130, 2000.
- McTiernan, J. M., Fisher, G. H., and Li, P.: The Solar Flare Soft X-Ray Differential Emission Measure and the Neupert Effect at Different Temperatures, *ApJ*, 514, 472–483, 1999.
- Pottasch, S.: On the Interpretation of the Solar Ultraviolet Emission Line Spectrum, *Space Sci. Rev.*, 3, 816–855, 1964.
- Reeves, K. K., Parenti, S., Reale, F., and Weber, M. A.: Methods of Analyzing Temperatures in Post-Flare Loops using the XRT on *Hinode*, AGU Fall Meeting 2007, SH51C-08, 2007.
- Schmelz, J. T., Scopes, J., Cirtain J. W., Winter H. D., and Allen J. D.: Observational Constraints on Coronal Heating Models Using Coronal Diagnostics Spectrometer and Soft X-Ray Telescope Data, *ApJ*, 556, 896–904, 2001.
- Sylwester, J., Schrijver, J., and Mewe, R.: Multitemperature analysis of solar X-ray line emission, *Solar Phys.*, 67, 285–309, 1980.
- Thomas, R. J., Starr, R., and Crannell, C. J.: Expressions to determine temperatures and emission measures for solar X-ray events from GOES measurements, *Solar Phys.*, 95, 323–329, 1985.
- Vaiana, G. S., Krieger, A. S., and Timothy, A. F.: Identification and Analysis of Structures in the Corona from X-Ray Photography, *Solar Phys.*, 32, 81–116, 1983.
- Weber, M. A., DeLuca, E. E., Golub, L., and Sette, A. L.: Temperature diagnostics with multichannel imaging telescopes, *IAU Symp.*, 223, 321–328, 2004.
- Withbroe, G. L.: The analysis of XUV emission lines, *Solar Phys.*, 45, 301–317, 1975.